

FULL PAPER

Open Access

Effect of lift force on the aerodynamics of dust grains in the protoplanetary disk

Masaki S Yamaguchi^{1*} and Shigeo S Kimura²

Abstract

We newly introduce *lift force* into the aerodynamics of dust grains in a protoplanetary disk. Although many authors have investigated the effects of the drag force, gravitational force, and electric force on the dust grains, the lift force has never been considered as a force exerted on dust grains in a gas disk. We show in this paper that the dust grains can be continuously spinning as a result of the frequent collisions and that the lift force continues to be exerted on them, which is valid in a certain parameter space where the grain size is larger than approximately 1 μm and where the distance from the central star is larger than 1 AU for the minimum mass solar nebula. In addition, we estimate the effects of the force on the grain motion and obtain results that show that the mean relative velocity between the grains due to the lift force is comparable to the gas velocity in the Kepler rotational frame when the Stokes number and lift-drag ratio are both approximately 1. This estimation is performed under the assumptions of steady state and the isotropic spin angular momentum. We also estimate the mean relative velocity when the grains keep spinning and conclude that the lift force marginally affects the mean relative velocity in the minimum mass solar nebula. If there is a grain-concentrated part in the disk, the relative velocity due to the lift force may dominate there because of the high collision rate.

Keywords: Gas disk; Protoplanetary; Aerodynamics; Dust grains

Background

In the theory of planet formation, the issue concerning the radial drift of meter-size dust remains an open question. In a typical scenario, micrometer-size dust grains grow to be kilometer-size planetesimals via collision and merging in a protoplanetary disk (Goldreich and Ward 1973; Hayashi 1981). When the dust grain grows to be meter-sized, there is velocity with respect to the disk gas and losses in the angular momentum due to the drag force. Thus, the grain falls down to the central star and cannot grow further (Adachi et al. 1976; Brauer et al. 2008).

Various scenarios have been proposed for solving the issue of meter-size dust. The gravitational instability in the dust layer was investigated first (Goldreich and Ward 1973; Sekiya 1983). In this scenario, the dust grains settle toward the mid-plane to form the dense layer, which then fragments into precursors of the planetesimals. However, the sedimentation of the grains leads to vertical shear

of the rotational velocity in the dust layer, which causes turbulence due to the Kelvin-Helmholtz instability. As a result, the grains cannot settle enough to form planetesimals (Sekiya 1998). The effects of turbulence due to magneto-rotational instability have also been considered (Balbus and Hawley 1976; Sano et al. 2000). Collisions between dust grains occur more frequently when there is an increase of the relative velocity due to the turbulence, so the growth rate of the grains also increases. Brauer et al. (2008), Okuzumi et al. (2012), and Kataoka et al. (2013) have taken into account these effects of the turbulence, and while the first study has found that the dust grains fall down to the central star if the grain density is relatively large, the second and third ones have found that the dust grains can grow at a sufficiently rapid rate to avoid the issue of meter-size dust when the grains are fluffy. As another scenario Youdin and Goodman (2005) have suggested, the planetesimals may be formed by the streaming instability caused by the interaction between the dust grains and disk gas.

We newly introduce *lift force* as a factor affecting the relative velocity between the dust grains. When a grain

*Correspondence: masaki.yamaguchi@nao.ac.jp

¹National Astronomical Observatory, 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan

Full list of author information is available at the end of the article

moves in fluid and when the fluid around the grain has circulation, the lift force is exerted on the grain perpendicularly to the grain velocity, and this is generally represented as

$$F_L = C_L \cdot \frac{\pi r_d^2}{m_d} \cdot \frac{1}{2} \rho_g |\vec{u}|^2, \quad (1)$$

where C_L , r_d , m_d , ρ_g , and \vec{u} represent the coefficient of the lift force, the radius and the mass of the grain, the density of the gas, and the velocity of a dust grain relative to the disk gas, respectively. We note here that F_L is defined as the lift force per unit mass. The coefficient of the lift force is determined by properties of the grain and flow. For a rotating sphere, the lift force is expressed as

$$\vec{F}_L = \frac{\pi \rho_g r_d^3}{m_d} \vec{\omega}_d \times \vec{u}, \quad (2)$$

where $\vec{\omega}_d$ represents the vector of the angular velocity of the grain. In Equation (2), Stokes law, which is valid when the Reynolds number is small, is adopted as the drag law (Rubinow and Keller 1961; Takagi 1974). In this case, the coefficient of the lift force is represented as

$$C_L = \frac{2m_d F_L}{\pi r_d^2 \rho_g u^2} = \frac{2r_d \omega_d}{u} \sin \theta, \quad (3)$$

where θ is the angle between $\vec{\omega}_d$ and \vec{u} . When the Reynolds number is so large that the turbulent flow is dominant, and when the Knudsen number is so large that the fluid cannot be regarded as a continuum, the lift force has not yet been formulated. Therefore, in this paper, we investigate the effects of the lift force only when Equation (2) can be applied. We estimate the conditions under which the lift force continues to be exerted on the grain in the ‘Sustainability of the spin of the dust grains’ section. We derive and reduce an equation of motion for the grain to estimate the relative velocity between the grains in the ‘Relative velocity between the dust grains’ section. In the ‘Discussion’ section, we evaluate the relative velocity between the grains when the lift force is kept exerted on the grains and discuss potential areas of improvement for our model. Finally, we summarize our study in the ‘Conclusions’ section.

Methods and results

Sustainability of the spin of the dust grains

In this section, we examine whether the dust grains keep spinning in the gas disk because the lift force does not act on non-spinning spherical objects. Here, we assume that the collisions between the dust grains induce the spin of the dust grains. The spinning dust is subjected to the torque due to the friction created by the background viscous fluid. After the spin-down timescale, the spin of the dust would stop. We can estimate the collision time t_{col} and the spin-down time $t_{\text{spin-down}}$. By comparing these

timescales, we obtain the parameter space where the lift force can act on the spinning dust grains. These timescales depend on the disk structure. We adopt the parameters for the disk structure in this paper as follows:

$$\Sigma_g = \Sigma_0 R_1^{-q}, \quad (4)$$

$$\Sigma_d = f_d \Sigma_g = f_d \Sigma_0 R_1^{-q}, \quad (5)$$

$$c_s = \sqrt{\frac{kT}{\bar{m}}} = c_{s,0} R_1^{-p}, \quad (6)$$

$$\Omega_K = \sqrt{\frac{GM_s}{R^3}} = \Omega_0 R_1^{-3/2}, \quad (7)$$

$$v_K = R\Omega_K = \Omega_0 R_0 R_1^{-1/2}, \quad (8)$$

$$H_g = \frac{\sqrt{2}c_s}{\Omega_K} = \frac{\sqrt{2}c_{s,0}}{\Omega_0} R_1^{-p+3/2}, \quad (9)$$

$$\rho_g = \frac{\Sigma_g}{\sqrt{\pi}H_g} = \frac{\Sigma_0 \Omega_0}{\sqrt{2\pi}c_{s,0}} R_1^{p-q-3/2}, \quad (10)$$

where R is the semi-major axis, R_0 is the typical radius of the disk, $R_1 = R/R_0$, f_d is the dust-to-gas mass ratio, $\bar{m} = 2.35m_H$ is the mean particle mass of gas, and $M_s = 1M_\odot$ is the mass of the central star. We use the isothermal sound speed c_s and the mid-plane gas density ρ_g when estimating timescales. If we choose $\Sigma_0 = 1.7 \times 10^3 \text{ g cm}^{-2}$, $c_{s,0} = 1.0 \times 10^5 \text{ cm s}^{-1}$, $\Omega_0 = 2.0 \times 10^{-7} \text{ s}^{-1}$, $R_0 = 1 \text{ AU}$, $f_d = 0.01$, $q = 3/2$, and $p = 1/4$, the disk profile is similar to the minimum mass solar nebula (MMSN; Hayashi 1981).

Collision timescale

The collision timescale is estimated as

$$t_{\text{col}} \sim (n_d \cdot \pi r_d^2 \cdot \langle v_{d-d} \rangle)^{-1}, \quad (11)$$

where n_d and v_{d-d} are the number density of the dust grains and the relative velocity between the dust grains, respectively. The parenthetic quantity $\langle Q \rangle$ represents the statistical average.

The dust number density is expressed as

$$n_d = \frac{\Sigma_d}{H_d m_d}, \quad (12)$$

where H_d is the scale height of the dust layer and m_d is the mass of the dust grains. We approximate that the mass distribution function of the dust is the delta function because it is necessary for the dust grains to collide with similar scale grains so that the grains gain the angular momentum. Considering the equilibrium between turbulent diffusion and sedimentation (Birnstiel et al. 2010), H_d is obtained as

$$H_d = H_g \cdot \left(\frac{\alpha}{\text{St}} \frac{1+2\text{St}}{1+\text{St}^2} \right)^{1/2}, \quad (13)$$

where $\text{St} \equiv \Omega_K t_s$ is the Stokes number (t_s is the stopping time by drag force). We use alpha prescription $\nu_{\text{turb}} = \alpha c_s H_g$ to describe the strength of the turbulence in the

protoplanetary disk and assume $\alpha < St$ to avoid the situation $H_g < H_d$. At the Stokes drag law regime, the Stokes number is written as

$$St = \frac{m_d \Omega_K}{6\pi r_d \rho_g \nu} = \frac{2\rho_{\text{int}} r_d^2 \Omega_K}{3\rho_g c_s \lambda_{\text{mfp}}} = \frac{2\sigma_{\text{mol}} \rho_{\text{int}} \Omega_0}{3\bar{m} c_{s,0}} r_d^2 R_1^{p-3/2}, \quad (14)$$

where ρ_{int} is the internal mass density of the dust grains, $\nu = c_s \lambda_{\text{mfp}}/3$ is the kinematic viscosity, and λ_{mfp} is the mean free path of the gas particles. The mean free path is estimated as $\lambda_{\text{mfp}} = \bar{m}/(\sigma_{\text{mol}} \rho_g)$, where σ_{mol} is the cross section of collisions between H_2 molecules. We adopt $\rho_{\text{int}} \simeq 3 \text{ g cm}^{-3}$ and $\sigma_{\text{mol}} \simeq 2 \times 10^{-15} \text{ cm}^2$. Equation (14) means that the Stokes number is independent of the normalization coefficient of the surface density Σ_0 , which is canceled out because of the one in λ_{mfp} .

Since the gas was assumed to be in a turbulent state described by the alpha prescription, we set the mean relative velocity $\langle v_{d-d} \rangle = \langle v_{d-d} \rangle_t$, where $\langle v_{d-d} \rangle_t$ means the relative velocity between the grains in the turbulent gas. According to (Ormel and Cuzzi 2007), $\langle v_{d-d} \rangle_t$ with similar scale grains can be represented as

$$\langle v_{d-d} \rangle_t = c_s \left(\frac{\alpha St}{\sqrt{1 + \frac{1}{4} St^2 (1 + St)^2}} \right)^{1/2}, \quad (15)$$

where we smoothly interpolate the two limiting solutions of $St \gg 1$ and $St \ll 1$. This expression is valid when the stopping time is larger than the turnover time of the Kolmogorov-scale eddy. The minimum size of the grain satisfying this condition is on the order of sub-millimeters for MMSN at 1 AU; hence, we focus on grains larger than approximately 1 mm in what follows.

Now we can express t_{col} as the function of r_d and R_1 by using Equations (11) to (15) as

$$t_{\text{col}} = \frac{4\sqrt{2}\rho_{\text{int}}}{3\Omega_0 \Sigma_0 f_d} r_d R_1^{q+3/2} f(St), \quad (16)$$

where

$$f(St) = \frac{1}{St} \left(\frac{1 + 2St}{1 + St^2} \right)^{1/2} \left(1 + \frac{1}{4} St^2 (1 + St)^2 \right)^{1/4}. \quad (17)$$

We note that when $\alpha < St$, t_{col} is independent of α because the effect of increasing $\langle v_{d-d} \rangle_t$ balances the one decreasing n_d . For the case with $St \ll 1$, the collision timescale is $t_{\text{col}} \propto St^{-1} r_d R_1^{q+3/2} \propto r_d^{-1} R_1^{3+q-p}$, while it is $t_{\text{col}} \propto St^{-1/2} r_d R_1^{q+3/2} \propto r_d^0 R_1^{2.25+q-0.5p}$ for the case with $St \gg 1$. If we adopt the same parameters as MMSN, the collision time is

$$t_{\text{col}} = 1.7 \times 10^6 f(St) R_1^3 r_{d,1} \text{ s}, \quad (18)$$

where $r_{d,1} = r_d/(1 \text{ cm})$.

Spin-down timescale

In the case of Stokes law, the angular momentum conservation around the spin axis of a spherical grain is given as

$$I_d \frac{d\omega_d}{dt} = -8\pi \rho_g \nu r_d^3 \omega_d, \quad (19)$$

where I_d is the moment of inertia of the grain. The torque acting onto a spherical body by viscous fluid is given in Rubinow and Keller (1961) and Takagi (1974). From this equation, the $t_{\text{spin-down}}$ is estimated as

$$t_{\text{spin-down}} = \frac{I_d}{8\pi \rho_g \nu r_d^3} = \frac{\rho_{\text{int}} r_d^2}{5\rho_g c_s \lambda_{\text{mfp}}} = \frac{\sigma_{\text{mol}} \rho_{\text{int}}}{5\bar{m} c_{s,0}} r_d^2 R_1^p. \quad (20)$$

In this second equation, we assume a spherical and uniform density grain whose moment of inertia is represented as $I_d = 2m_d r_d^2/5$. The spin-down time becomes longer as the dust grain becomes larger. Here, we note that the spin-down time is independent of Σ_0 for the same reason as the Stokes number [see Equation (14)]. For MMSN, $t_{\text{spin-down}}$ is estimated as

$$t_{\text{spin-down}} = 3.0 \times 10^3 r_{d,1}^2 R_1^{1/4} \text{ s}. \quad (21)$$

Comparison of timescales

Now, we can obtain the size of dust grains that are able to keep spinning. We estimate these timescales just in the Stokes law regime because the lift force in other regimes is uncertain. There are two necessary conditions to realize Stokes law. One is that the gas can be regarded as a continuum medium, which is expressed as $r_d \gtrsim 9\lambda_{\text{mfp}}/4$. The other is that the flow around the dust grains is laminar, which is represented as $Re = 2ur_d/\nu \lesssim 20$ (Shirayama 1992), where u is the relative velocity of the dust to the gas. Here, we should actually include the effect of turbulence in the expression of u as in Ormel and Cuzzi (2007) so that the physical situation is consistent with that of Equation (15). However, taking this effect into account causes complicated equations. Thus, as a first-step attempt, we assume that u is equal to the relative velocity between the orbital velocity of the gas and the Keplerian velocity, i.e., $u = \eta v_K$, where

$$\eta \equiv \frac{2p + 2q + 3}{4} \left(\frac{c_s}{v_K} \right)^2, \quad (22)$$

which is given in (Adachi et al. 1976). By these conditions, we find that our estimation is valid in the following range:

$$\begin{aligned} r_{d,\text{min}} &\equiv \frac{9\bar{m}c_{s,0}}{2\sigma_{\text{mol}} \Sigma_0 \Omega_0} R_1^{q-p+3/2} \lesssim r_d \lesssim r_{d,\text{max}} \\ &\equiv \frac{40\sqrt{2\pi}\bar{m}R_0}{3(2p + 2q + 3)\sigma_{\text{mol}} \Sigma_0} R_1^{q+1}. \end{aligned} \quad (23)$$

For MMSN, this condition is simply written as

$$3.2R_1^{11/4} \lesssim r_{d,1} \lesssim 89R_1^{5/2}. \quad (24)$$

Figure 1 shows the two timescales t_{col} (solid lines) and $t_{\text{spin-down}}$ (dashed lines) at $R_1 = 1$ for MMSN. We plotted t_{col} and $t_{\text{spin-down}}$ in the range that satisfies the condition (24). From Figure 1, we can see that t_{col} is larger than $t_{\text{spin-down}}$, so that the spin of the dust would stop at $R_1 = 1$. The difference between the two timescales gets smaller as the dust grains become larger. Equations (18) and (21) show that large grains are likely to satisfy the condition $t_{\text{col}} < t_{\text{spin-down}}$. From Equation (24), the Stokes regime can be adopted for the larger dust grains at the outer region of the disk. Thus, we expect that the condition $t_{\text{col}} < t_{\text{spin-down}}$ is satisfied at the outer region $R_1 > 1$. Figure 2 shows the parameter space where the dust grains keep spinning in the R_1 - r_d plane for MMSN. The Stokes regime is realized between the dashed green lines. The condition $t_{\text{col}} < t_{\text{spin-down}}$ is satisfied above the solid red line. In the blue region, the condition is satisfied with the Stokes regime. There are dust grains that keep spinning with the Stokes regime in $R_1 \gtrsim 1.3$. The grains that can keep spinning have the size $r_d \sim r_{d,\text{max}}$. The dotted magenta line shows the dust radius when $\text{St} = 1$, which is used in the ‘Discussion’ section.

Relative velocity between the dust grains

In this section, we investigate whether the mean relative velocity is comparable to or greater than the gas velocity

in the Kepler rotational frame. Since this gas velocity is comparable to the typical relative velocity between a large grain and a small one compared to 1-m-sized dust, we take it as a reference value. First, we derive the equation of motion for a dust grain assuming that it moves at a terminal velocity. Next, we estimate the mean relative velocity by assuming an isotropic distribution for the spin angular momentum.

Here, for simplicity, we assume that the dust grains move on the mid-plane of the disk, which means that the z -component of the lift force is assumed to be zero, where the z -axis is taken as the disk axis, and we adopt the cylindrical coordinates described below. Since the direction of the spin angular momentum can be taken arbitrarily, the lift force can show the z -component. Nevertheless, we neglect the z -component of the velocity to simplify the calculation below.

For preparation to derive the equation of motion, we express a projected vector of the lift force on the mid-plane in terms of the direction of the spin angular momentum of a dust grain. Since the direction of the lift force is perpendicular to the spin angular momentum and the velocity of the grain with respect to the gas, then

$$\vec{F}_L = A\vec{\omega}_d \times \vec{u}, \quad (25)$$

where the coefficient satisfies $A = \frac{\pi\rho_g r_d^3}{m_d}$ (see the ‘Background’ section). Since the z -component of \vec{u} is zero,

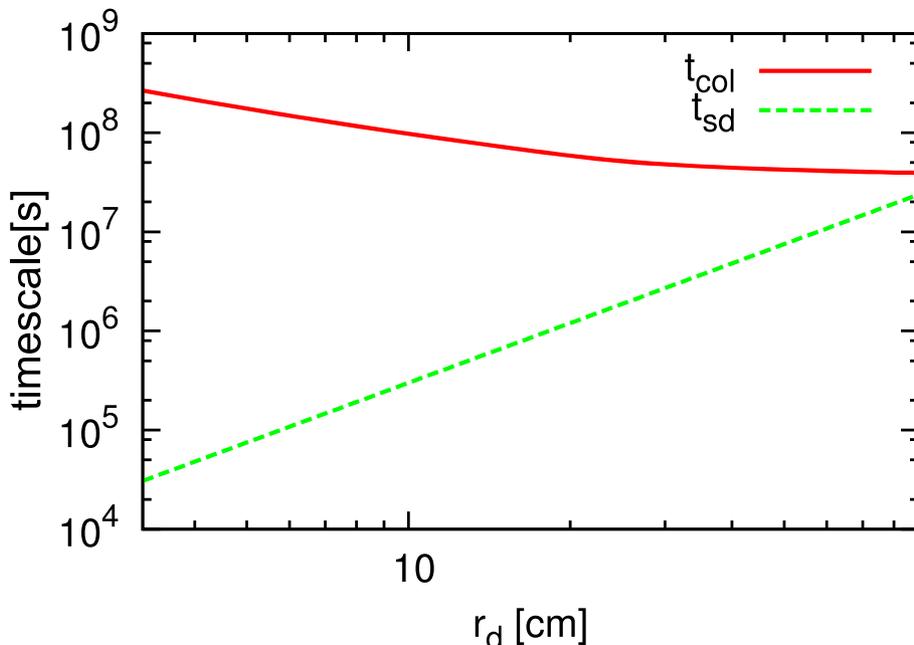


Figure 1 r_d dependence of t_{col} (solid lines) and $t_{\text{spin-down}}$ (dashed lines) at $R_1 = 1$. t_{col} is always longer than $t_{\text{spin-down}}$ in the Stokes regime.

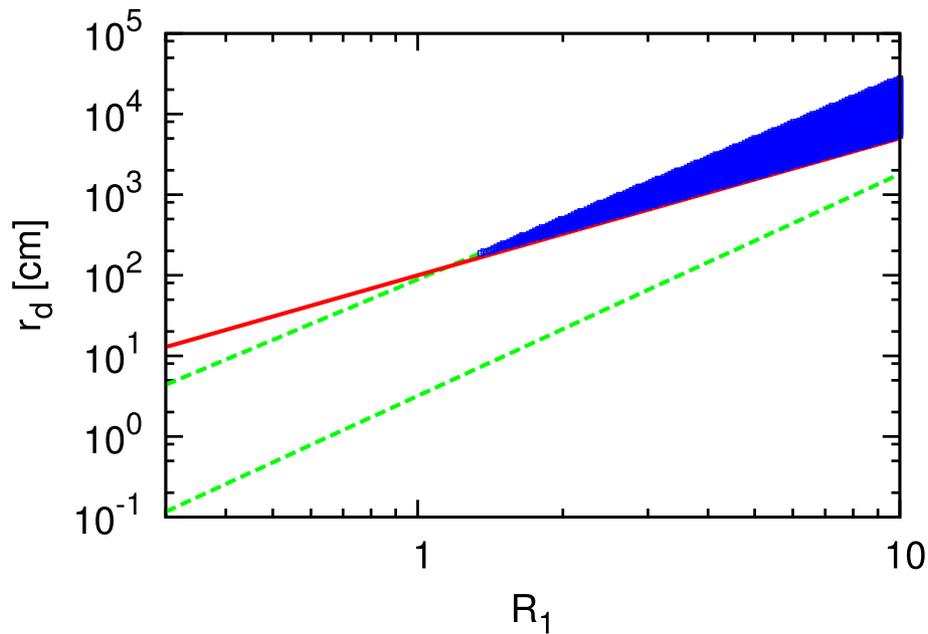


Figure 2 Parameter space where the dust grains can keep spinning in the R_1 - r_d plane. The Stokes regime is between the two dashed green lines. The condition $t_{\text{col}} < t_{\text{spin-down}}$ is satisfied above the solid red line. The grains may keep spinning in the blue region. The dotted magenta line represents the line where the condition $\text{St} = 1$ is satisfied.

the lift force vector projected on the mid-plane $\vec{F}_{L,\text{mid}}$ is expressed as

$$\vec{F}_{L,\text{mid}} = \vec{F}_L \cdot \begin{pmatrix} \vec{e}_r \\ \vec{e}_\theta \end{pmatrix} = A\omega_d\mu \begin{pmatrix} -u_\theta \\ u_r \end{pmatrix} \equiv F_L \begin{pmatrix} -u_\theta/u \\ u_r/u \end{pmatrix}, \quad (26)$$

where \vec{e} with a subscript and μ represent a unit vector in the direction of the subscript and the cosine of the angle between $\vec{\omega}_d$ and the z -axis, respectively. We note here that the F_L values do not depend on the azimuth angle of the spin angular momentum.

Next, we derive and reduce the equation of motion for a dust grain. Now, the forces exerted on the dust grain are the gravitational force of the central star, the drag force, and the lift force, so the equation of motion is expressed as

$$\frac{d\vec{v}}{dt} = -\frac{GM_\odot}{r^2}\vec{e}_r - \vec{F}_D + \vec{F}_{L,\text{mid}}, \quad (27)$$

where we assume that the mass of the central star is the same as the solar one. As the first step for the reduction of Equation (27), we divide it into two equations for the r and θ components. Since the velocity of the disk gas does not have the radial component, the components of the velocity of the dust grain are represented as $(u_r, u_\theta) = (v_r, v_\theta - r\Omega_g)$, where Ω_g is the orbital angular velocity of

the disk gas around the central star. Thus, Equation (27) is expressed as

$$\frac{dv_r}{dt} - \frac{v_\theta^2}{r} = -\frac{GM_\odot}{r^2} - F_D \frac{v_r}{u} - F_L \frac{v_\theta - r\Omega_g}{u}, \quad (28)$$

$$\frac{dv_\theta}{dt} + \frac{v_r v_\theta}{r} = -F_D \frac{v_\theta - r\Omega_g}{u} + F_L \frac{v_r}{u}. \quad (29)$$

As the second step, we transform this into coordinates rotating at the angular velocity of the Kepler rotation, that is, $v_\theta = v_K + v'_\theta$. As the third step, we assume that the motion of the dust grain is stationary and that $|v_r|, |v'_\theta| \ll v_K$.

This stationary assumption may be invalid when taking into account the timescales discussed in the ‘Sustainability of the spin of the dust grains’ section. The stopping time t_s is represented as

$$t_s = \frac{m_d}{6\pi r_d \rho_g v} \sim 10^4 r_{d,1}^2 R_1^{1/4} \sim 4 t_{\text{spin-down}}. \quad (30)$$

This means that the dust grain stops spinning before moving at the terminal velocity independently of the dust size and the distance from the central star. Thus, as long as the lift force is exerted on the grain, the motion of the grain cannot reach a steady state. Alternatively, the grain motion can be considered to be determined by the merger of the parent grains (or scattering by the other grains). Nevertheless, we assume that the grain motion reaches the steady state for the first stage of this type of work.

Finally, we non-dimensionalize the variables as

$$x = \frac{v_r}{\eta v_K}, y = \frac{v'_\theta}{\eta v_K}, g_D = \frac{F_D}{u\Omega_K}, g_L = \frac{F_L}{u\Omega_K}, \quad (31)$$

where η is the constant satisfying the equation $r\Omega_g = v_K(1 - \eta)$ and Ω_K is the angular velocity of the Kepler motion. Thus, we obtain two algebraic equations:

$$2y = g_D x + g_L (y + 1), \quad (32)$$

$$\frac{1}{2}x = -g_D (y + 1) + g_L x. \quad (33)$$

These equations represent the balance between the Coriolis, drag, and lift forces.

The solution of the equations is

$$x = \frac{-2g_D}{g_D^2 + (g_L - 2)(g_L - 1/2)}, \quad (34)$$

$$y = -1 + \frac{-2(g_L - 1/2)}{g_D^2 + (g_L - 2)(g_L - 1/2)}. \quad (35)$$

Here, Equations (26) and (31) lead to $g_L = g_{L,max}\mu$, where $g_{L,max} \equiv \frac{A\omega_d}{\Omega_g}$, and we introduce the lift-drag ratio $R_{LD} \equiv g_{L,max}/g_D$ to obtain $g_L = g_D R_{LD}\mu$. Thus, x and y are expressed as functions of μ , g_D , and R_{LD} . We show $x(\mu)$ and $y(\mu)$ in Figure 3. Here, we assume $g_D = 1$ and $R_{LD} = 1$ as trial values.

Since $x = -1$ when we neglect the lift force, the curve of $x(\mu)$ shows that the radial velocity of the dust grain can be a third or four times of that without the lift force. On the other hand, when $\mu < 0$, y is almost constant and comparable to that without the lift force. When μ is larger than 0.5, y is smaller than -1 , which means that the dust grain orbits more slowly than the gas. In addition, we see that x and y decrease as μ is close to unity. Therefore, the absolute value of the velocity tends to increase as μ increases.

Next, we calculate the average and dispersion of the velocity of the dust grain on the disk mid-plane, assuming that the spin angular momentum is isotropic, which is just for simplicity. Thus, the direction distribution satisfies $f(\Omega) = \frac{1}{4\pi}$, which is equivalent to $f(\mu) = \frac{1}{2}$, where Ω is a solid angle parameter. The average and dispersion of x are calculated by performing the integration below.

$$\begin{aligned} \langle x \rangle &= \int x f(x) dx \\ &= \frac{1}{2} \int_{-1}^1 x(\mu) d\mu, \\ \langle x^2 \rangle &= \int x^2 f(x) dx \\ &= \frac{1}{2} \int_{-1}^1 x^2(\mu) d\mu, \end{aligned} \quad (36)$$

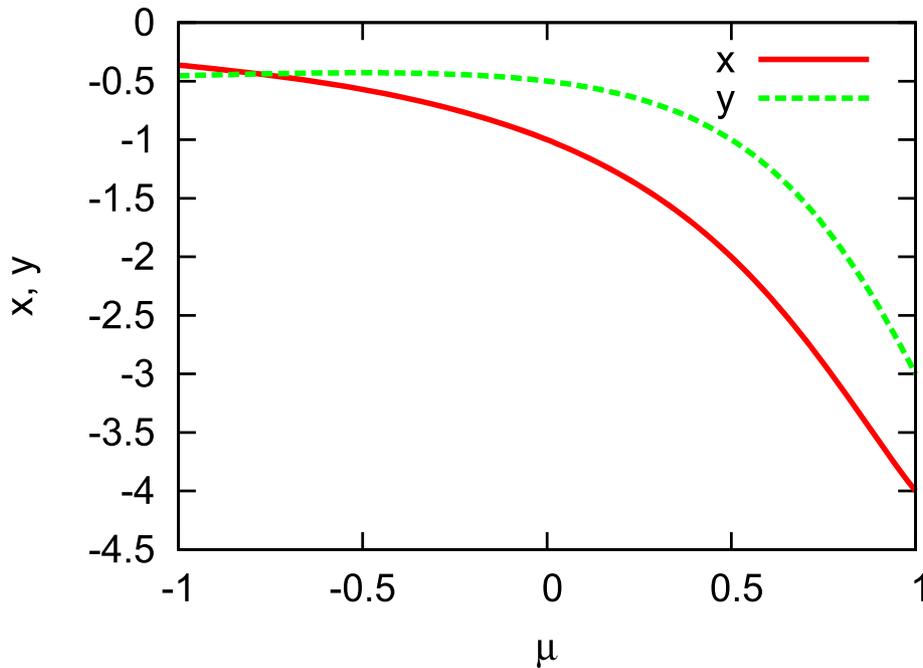


Figure 3 Dependence of the grain velocity on the direction of the spin angular momentum. We show the grain velocity in the radial direction x (the solid line) and in the azimuthal direction y (the dashed line). The variable μ represents the cosine of the angle between the spin angular momentum and z -axis. We take the parameters as $g_D = 1$ and $R_{LD} = 1$.

where we transform the integration variable into μ . We also can derive the same expression for y ,

$$\langle y \rangle = \frac{1}{2} \int_{-1}^1 y(\mu) d\mu, \quad (37)$$

$$\langle y^2 \rangle = \frac{1}{2} \int_{-1}^1 y^2(\mu) d\mu.$$

By taking $g_D = 1$ and $R_{LD} = 1$, we can obtain the approximate value of the average and standard deviation of the velocity,

$$\langle x \rangle \simeq -1.4, \quad (38)$$

$$\langle y \rangle \simeq -1.2, \quad (39)$$

$$\sqrt{\langle x^2 \rangle} \simeq 1.0, \quad (40)$$

$$\sqrt{\langle y^2 \rangle} \simeq 0.6, \quad (41)$$

$$\langle w_{rel} \rangle \equiv \sqrt{\langle x^2 \rangle + \langle y^2 \rangle} \simeq 1.2. \quad (42)$$

Equation (38) means that the dust grains, on average, fall down to the star faster than without the lift force. Moreover, Equation (39) means that their average orbit is at almost the same velocity as the gas. We note that the standard deviation of the velocity represents the average of the relative velocity between the grains. Therefore, Equation (42) means that the average relative velocity exceeds the relative velocity between the gas and Kepler velocity, so that the collision rate is affected by the lift force when $g_D = 1$ and $R_{LD} = 1$.

We finally calculate the averaged relative velocity on the disk mid-plane $\langle w_{rel} \rangle$ for arbitrary values of g_D and R_{LD} . Figure 4 shows the contour lines of $\langle w_{rel} \rangle = 1$ (the solid line) and 0.1 (the dashed line) on the g_D - R_{LD} plane. We see that the relative velocity is large when g_D is small and when R_{LD} is large, which corresponds to the situation where the lift force is efficiently exerted on the grains. The important fact is that there exists a region satisfying $\langle w_{rel} \rangle \gg 1$, where the lift force non-negligibly affects the dynamics of the system of grains, compared to the case without the force.

Discussion

Lift-drag ratio

In this subsection, we estimate the lift-drag ratio R_{LD} to investigate how efficiently the lift force affects the motion of the dust. In the Stokes regime, the lift coefficient C_L and drag coefficient C_D are represented as

$$C_D = \frac{4c_s \lambda_{mfp}}{r_d u}, \quad (43)$$

$$C_L = \frac{2m_d F_L}{\pi r_d^2 \rho_g u^2} = \frac{2r_d \omega_d}{u} \sin \theta. \quad (44)$$

The lift coefficient depends on the spin angular velocity ω_d of the dust. Here, we estimate ω_d induced by the

collisions of the grains. When two grains with the same mass m_d collide with the impact parameter b , the angular momentum around the center of mass is represented by

$$L = \frac{bv_{d-d}m_d}{2}. \quad (45)$$

Given the weight by a cross section, we derive the averaged angular velocity $\sqrt{\langle L^2 \rangle}$ as

$$\sqrt{\langle L^2 \rangle} = \left(\frac{\int_0^{2r_d} L^2 2\pi b db}{\int_0^{2r_d} 2\pi b db} \right)^{1/2} = \frac{1}{\sqrt{2}} \langle v_{d-d} \rangle r_d m_d. \quad (46)$$

If we assume that the grain obtains the mass $2m_d$ and this angular momentum after the collision, the resultant angular velocity ω_d is represented as

$$\omega_d = \frac{5\sqrt{2}}{4} \frac{\langle v_{d-d} \rangle}{r_d}. \quad (47)$$

Using Equations (43), (44), and (47), we can reduce the lift-drag ratio R_{LD} as

$$R_{LD} = \frac{C_L}{C_D} = \frac{5\sqrt{2}r_d \langle v_{d-d} \rangle}{8c_s \lambda_{mfp}}. \quad (48)$$

Furthermore, we adopt $\langle v_{d-d} \rangle = \langle v_{d-d} \rangle_t$, so that this expression is nearly independent of the dust radius r_d for $St > 1$. For $St \gg 1$, the lift-drag ratio approaches asymptotically to a maximum,

$$R_{LD} \simeq 20\alpha^{1/2} R_1^{-17/8}, \quad (49)$$

where we take the MMSN disk parameters.

We see from Figure 2 that when r_d is approximately 10^2 cm at $R_1 \sim 1$, which corresponds to $g_D = St^{-1} \sim 0.1$, then the conditions where $t_{col} \sim t_{spin-down}$ and the drag force is represented with Stokes law are marginally satisfied. In this case, we find R_{LD} of approximately 2.0 when $\alpha = 0.01$ using Equation (49). Thus, from Figure 4, we obtain $\langle w_{rel} \rangle \sim 0.1$, which means that the averaged relative velocity due to the lift force is a tenth of ηv_K for the 1- μ m-sized dust at 1 AU from the central star.

Dependence of the relative velocity on disk parameters

The situation stated in the previous subsection can be qualitatively or quantitatively changed by adopting disk parameters different from MMSN. If we take an f_d larger than 0.01, as proposed in Sekiya (1998) and Hasegawa and Tsuribe (2014), the averaged relative velocity $\langle w_{rel} \rangle$ can be approximately 1. The dust-to-gas ratio f_d is included just in the expression of t_{col} [Equation (16)], so that a larger f_d means a smaller t_{col} and thereby smaller r_d satisfying $t_{col} = t_{spin-down}$. This implies that the red line in Figure 2 moves down and that the blue region expands inside. When f_d is approximately 0.1, the blue region includes the dotted magenta line, which shows the grain

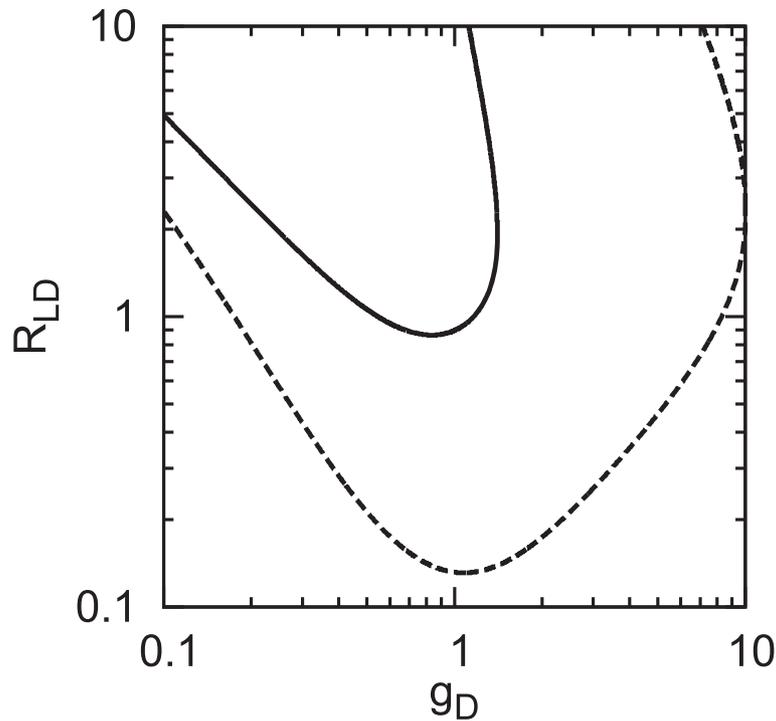


Figure 4 Contours of the averaged relative velocity of the grains are shown on the g_D - R_{LD} plane. The solid and dashed lines represent $\langle w_{rel} \rangle = 1$ and 0.1, respectively.

size satisfying $St = 1$ at $R_1 \sim 0.5$. Since R_{LD} is approximately 1 when $\alpha = 10^{-4}$, we find from Figure 4 that the relative velocity $\langle w_{rel} \rangle$ is approximately 1. In this case, the relative velocity due to the lift force $\eta v_K \langle w_{rel} \rangle$ is approximately 60 m s^{-1} , where we note that ηv_K does not depend on R_1 when $p = 1/4$. This value exceeds the relative velocity due to the turbulence $\langle v_{d-d} \rangle \sim 8 \text{ m s}^{-1}$, which is computed with Equation (15). Therefore, the lift force is expected to efficiently affect the growth rate of the grain whose size is 10 cm at $R_1 \sim 0.5$ if there is a grain-concentrated part.

Recently, a steeper density profile, which is denser at 1 AU, has been proposed (Desch 2007). When the density profile is steeper, a larger q is adopted. If q is larger, the radial profile of the minimum and maximum sizes of the grains is steeper [Equation (23) and dashed green lines in Figure 2]. Moreover, we can find that the red line in Figure 2 does not change as much as the green lines by comparing Equation (16) to (20). Thus, the blue region in Figure 2 slightly shifts to the inner region. Therefore, we expect that when the density profile is steeper with the normalization coefficient fixed, the innermost radius where the lift force continues to be exerted on the grains slightly decreases close to $R_1 = 1$.

Suppose that the disk has a larger surface density. Spin-down time [Equation (20)] and Stokes number [Equation (14)] are independent of the surface density Σ_0 .

Conversely, collision time is inversely proportional to Σ_0 [Equation (16)], so that r_d satisfying $t_{col} = t_{spin-down}$ is proportional to $\Sigma_0^{-1/2}$ for $St \gg 1$. Additionally, $r_{d,min}$ and $r_{d,max}$ in Equation (23) are inversely proportional to Σ_0 . Therefore, the blue region moves outside and the minimum grain size in the blue region is nearly unchanged. The ratio R_{LD} becomes smaller with decreasing R_1 , so that the relative velocity due to the lift force is smaller than that for the fiducial surface density.

Model refinement for other effects

We can refine the model in this paper by taking into account realistic porosity and shape of the dust grain (Suyama et al. 2008). If the grain is fluffy, ρ_{int} is smaller than the value we used in the ‘Sustainability of the spin of the dust grains’ section. Equations (14), (18), and (17) lead to the dependence of t_{col} on ρ_{int} ,

$$t_{col} \propto \begin{cases} \rho_{int}^{\frac{1}{2}} & \text{for } St \gg 1 \\ \rho_{int}^0 & \text{for } St \ll 1, \end{cases} \quad (50)$$

while $t_{spin-down} \propto \rho_{int}^1$. Thus, when ρ_{int} decreases, $t_{spin-down}$ decreases more rapidly than t_{col} , which means that the lift force is exerted on the grain over a shorter period of time. If the grain is lumpy, the coefficient of lift would become as large as a baseball or a golf ball.

It is also worth taking into account realistic collision processes between the grains, such as simple scattering (bouncing), minor merger, and destruction. Through the simple scattering, the grain may gain the spin angular momentum by means of surface friction, where the maximum surface velocity of the scattered grains is v_{d-d} , which is less than the mean surface velocity in the case of the major merger [Equation (47)]. In addition, in the case of a minor merger, which is realized when we consider the size distribution of grains (Windmark et al. 2012), the grains obtain less spin angular momentum compared to the case of the major merger. Thus, we expect less mean relative velocity when the grains undergo the scattering and minor merger. On the other hand, if the destruction (fragmentation) occurs when the grains collide, the grains may gain larger spin angular momentum. The experiments show that agglomerates (centimeter to decimeter size) are divided into many fragments that are of millimeter to centimeter size through low-velocity collisions (Deckers and Teisere 2012; Schr apler et al. 2012). If these fragments have a large spin angular momentum, they can be sufficiently affected by the lift force. How much spin angular momentum grains gain depends on many parameters, but we will defer this question to future work.

The z -component of the lift force, which is omitted in this paper, may affect the resultant relative velocity of the dust grain. The equation of motion in the z -direction includes the gravitational force by the central star, the drag force, and the lift force. Although the steady state cannot be realized because the gravitational force depends on the altitude from the disk mid-plane, the dust grains should gain a momentum in the z -direction. This causes an increase in the absolute value of the velocity, which may result in an increase in the relative velocity.

Conclusions

In this paper, we investigated the effects of the lift force on dust grains in a protoplanetary disk from two perspectives. We first investigated whether the lift force is kept exerted on the grains or not. We assumed the grains are in the minimum mass solar nebula where the turbulence develops. We estimated the collision timescale and the spin-down timescale and found that the grain keeps spinning as a result of collisions with the other grains if the radius of the grain is larger than 100 cm at $\gtrsim 1$ AU from the central star.

Next, we calculated the mean relative velocity between the grains caused by the lift force. The grains obtain spin angular momenta with various directions by collisions between themselves, so that the lift forces exerted on them have various directions as well. Thus, the relative velocity yields between the grains. We assumed that the grains are in the steady state and that the distribution of their spin momenta shows the isotropy. We then showed that

the mean relative velocity is comparable to the gas velocity at the Kepler rotational frame when $F_L \gtrsim F_D$ and $t_s \approx 1/\Omega_K$, where F_L , F_D , t_s , and Ω_K are the lift force, the drag force, the stopping time of the grains by the drag, and the Kepler angular velocity, respectively. This means that the lift force can sufficiently affect the collision rate, which affects the growth rate of the grains under the parameter set.

We also estimated the mean relative velocity when the grains keep spinning by combining the above two results. We found that for the minimum mass solar nebula, the mean relative velocity due to the lift force is smaller than the gas velocity at the Kepler rotational frame. We present the mean relative velocity as being comparable to the gas velocity if the disk has grain-concentrated parts where the dust-gas ratio is ten times larger than MMSN, so the lift force may affect the collision rate in the parts.

Abbreviations

MMSN: Minimum mass solar nebula.

Competing interests

The authors declare that they have no competing interests.

Authors' contributions

MSY carried out the estimation of the mean relative velocity of the grains, participated in the sequence alignment, and drafted the manuscript. SSK carried out the estimation of the condition under which the lift force is kept exerted on the grains, participated in the sequence alignment, and drafted the manuscript. Both authors read and approved the final manuscript.

Acknowledgements

MSY and SSK thank T. Tsuribe, S. Inutsuka, and S. Okuzumi for useful discussions and comments and acknowledge the anonymous referees for useful comments as well. SSK is supported by a Grant-in-Aid from the JSPS (Japan Society for the Promotion of Science) Research Fellowships for Young Scientists program (A2517840).

Author details

¹National Astronomical Observatory, 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan. ²Department of Earth and Space Science, Graduate School of Science, Osaka University, 1-1 Machikaneyama, Toyonaka, Osaka 560-0043, Japan.

Received: 30 April 2014 Accepted: 23 September 2014

Published: 3 October 2014

References

- Adachi I, Hayashi C, Nakazawa K (1976) The gas drag effect on the elliptical motion of a solid body in the primordial solar nebula. *Prog Theor Phys* 56:1756–1771
- Balbus SA, Hawley JF (1976) A powerful local shear instability in weakly magnetized disks. I - Linear analysis. II - Nonlinear evolution. *Astrophys J* 376:214–233
- Birnstiel T, Dullemond CP, Brauer F (2010) Gas- and dust evolution in protoplanetary disks. *Astron Astrophys* 513:79–99
- Brauer F, Dullemond CP, Henning Th (2008) Planetesimal formation near the snow line in MRI-driven turbulent protoplanetary disks. *Astron Astrophys* 487:1–4
- Deckers J, Teiser J (2012) Colliding decimeter dust. *Astrophys J* 769:151–156
- Desch SJ (2007) Mass distribution and planet formation in the solar nebula. *Astrophys J* 671:878–893
- Goldreich P, Ward W (1973) The formation of planetesimals. *Astrophys J* 183:1051–1062
- Hasegawa Y, Tsuribe T (2014) Kelvin–Helmholtz instabilities in multi-sized dust layers. *Publ Astron Soc Jpn* 66:L2

- Hayashi C (1981) Structure of the solar nebula, growth and decay of magnetic fields and effects of magnetic and turbulent viscosities on the nebula. *Prog Theor Phys Suppl* 70:35–53
- Kataoka A, Tanaka H, Okuzumi S, Wada K (2013) Fluffy dust forms icy planetesimals by static compression. *Astron Astrophys* 557:L4
- Okuzumi S, Tanaka H, Kobayashi H, Wada K (2012) Rapid coagulation of porous dust aggregates outside the snow line: a pathway to successful icy planetesimal formation. *Astrophys J* 752:106–123
- Ormel CW, Cuzzi JN (2007) Closed-form expressions for particle relative velocities induced by turbulence. *Astron Astrophys* 466:413–420
- Rubinow SI, Keller JB (1961) The transverse force on a spinning sphere moving in a viscous fluid. *J Fluid Mech* 11:447–459
- Sano T, Miyama SM, Umebayashi T, Nakano T (2000) Magnetorotational instability in protoplanetary disks. II. Ionization state and unstable regions. *Astrophys J* 543:486–501
- Schräpler R, Blum J, Seizinger A, Kley W (2012) The physics of protoplanetary dust agglomerates. VII. The low-velocity collision behavior of large dust agglomerates. *Astrophys J* 758:35–43
- Sekiya M (1983) Gravitational instabilities in a dust-gas layer and formation of planetesimals in the solar nebula. *Prog Theor Phys* 69:1116–1130
- Sekiya M (1998) Quasi-equilibrium density distributions of small dust aggregations in the solar nebula. *Icarus* 133:298–309
- Shirayama S (1992) Flow past a sphere: topological transitions of the vorticity field. *AIAA J* 30:349–358
- Suyama T, Wada K, Tanaka H (2008) Numerical simulation of density evolution of dust aggregates in protoplanetary disks. I. Head-on collisions. *Astrophys J* 684:1310–1322
- Takagi H (1974) On the slow motion of a sphere in a viscous fluid. *J Phys Soc Jpn* 37:505–510
- Windmark F, Birnstiel T, Güttler C, Blum J, Dullemond CP, Henning Th (2012) Planetesimal formation by sweep-up: how the bouncing barrier can be beneficial to growth. *Astron Astrophys* 540:73–89
- Youdin AN, Goodman J (2005) Streaming instabilities in protoplanetary disks. *Astrophys J* 620:459–469

doi:10.1186/1880-5981-66-132

Cite this article as: Yamaguchi and Kimura: **Effect of lift force on the aerodynamics of dust grains in the protoplanetary disk.** *Earth, Planets and Space* 2014 **66**:132.

Submit your manuscript to a SpringerOpen® journal and benefit from:

- Convenient online submission
- Rigorous peer review
- Immediate publication on acceptance
- Open access: articles freely available online
- High visibility within the field
- Retaining the copyright to your article

Submit your next manuscript at ► springeropen.com
